

# THE RADIO AFTERGLOW FROM THE GIANT FLARE OF SGR 1900+14: THE SAME MECHANISM AS AFTERGLOWS FROM CLASSIC GAMMA-RAY BURSTS?

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## ABSTRACT

A radio afterglow was detected following the 1998 August 27 giant flare from the soft gamma repeater (SGR) 1900+14. Its short-lived behavior is quite different from the radio nebula of SGR 1806-20, but very similar to radio afterglows from classic gamma-ray bursts (GRBs). Motivated by this, we attempt to explain it with the external shock model as invoked in the standard theory of GRB afterglows. We find that the light curve of this radio afterglow is not consistent with the forward shock emission of an ultra-relativistic outflow, which is suggested to be responsible for the initial hard spike of the giant flare. Nevertheless, shock emission from a mildly or sub-relativistic outflow expanding into the interstellar medium could fit the observations. The possible origin for this kind of outflow is discussed, based on the magnetar model for SGRs. Furthermore, we suggest that the presence of an ultra-relativistic fireball from SGR giant flares could be tested by rapid radio to optical follow-up observations in the future.

*Subject headings:* gamma rays: bursts—stars: individual (SGR 1900+14)—ISM: jets and outflows

## 1. INTRODUCTION

Soft gamma repeaters (SGRs) are generally characterized by sporadic and short ( $\sim 0.1$ s) bursts of hard X-rays with luminosities as high as  $10^4$  Eddington luminosity. They are also well-known for two giant flares: the first on 1979 March 5 from SGR 0526-66 (Mazets et al. 1979) and the second on 1998 August 27 from SGR 1900+14 (Hurley et al. 1999). Frail et al. (1999) reported, following the giant August 1998 flare from SGR 1900+14, the detection of a transient radio source. Their observations covered the time from about one week to one month after the flare. The spectrum between 1 and 10 GHz is well fitted by a power-law with  $F_\nu \propto \nu^{-0.74 \pm 0.15}$ . The source appears to have peaked at about a week after the burst and subsequently undergone a power law decay with an exponent of  $\alpha = -2.6 \pm 1.5$ .

The initial hard spike of the August 27 flare has a duration of  $\sim 0.5$ s and luminosity greater than  $2 \times 10^{44}$  ergs<sup>-1</sup> ( $>15$ KeV) if the source distance is  $d \simeq 7$ Kpc (Vasisht et al. 1994). The short duration, high luminosity and hard spectrum indicate that a relativistically expanding fireball was driven from the star. The fireball should be relatively clean and the Lorentz factor  $\Gamma \gtrsim 10$  was inferred from the luminosity and the temporal structure (Thompson & Duncan 2001). With the experience of GRB afterglows, one may naturally ask whether this power-law fading radio afterglow is due to the blast wave emission driven by the fireball. Huang, Dai & Lu (1998) and Eichler (2003) had made some discussions on the possible afterglow emission from SGRs. In this paper, we try to explain the radio afterglow from this flare. We study the afterglow emission from the giant flare in section 2. We find that shock emission from an ultra-relativistic outflow fails to explain this radio afterglow. However, we propose that a mildly or sub-relativistic outflow expanding into the interstellar medium could fit the observations. Finally, we discuss the possible origin for this kind of outflow in section 3.

## 2. RADIO AFTERGLOW FROM SGR GIANT FLARES

We consider that an outflow with “isotropic” kinetic energy  $E_0$  and Lorentz factor  $\Gamma_0$  ejected from the SGR expands into the ambient medium with a constant number density  $n$ . The interaction between the outflow and the surrounding medium is analogous to GRB external shock (Rees & Mészáros 1992; Mészáros & Rees 1997), but with quite different  $E_0$  and  $\Gamma_0$ . The Sedov time at which the shock enters the non-relativistic phase is roughly given by  $t_{\text{nr}} = (3E_0/4\pi n m_p c^5)^{1/3} = 1\text{days} (E_{0,44}/n_0)^{1/3}$ , where  $m_p$  is the proton mass and we used the usual notation  $a \equiv 10^n a_n$ . As the shock must have entered the non-relativistic phase during the observation time of the radio afterglow from the giant flare, we develop a model which holds for both the relativistic and non-relativistic phases. From the view of the energy conservation, the dynamic equation can be approximately simplified as (e.g. Huang, Dai & Lu 1999; Wang, Dai & Lu 2003)

$$(\Gamma - 1)M_0 c^2 + (\Gamma^2 - 1)m_{\text{sw}} c^2 = E_0, \quad (1)$$

where  $\Gamma$  is the Lorentz factors of the outflow,  $m_{\text{sw}} = (4/3)\pi R^3 m_p n$  is the mass of the swept-up ISM ( $R$  is the shock radius) and  $M_0$  is the mass of the original outflow. The first term on the left of the equation is the kinetic energy of the outflow and the second term is the internal energy of the shock.

The kinematic equation of the ejecta is

$$dR/dt = \beta c / (1 - \beta), \quad (2)$$

where  $v = \beta c$  is the bulk velocity of the outflow with  $\beta(\Gamma) = (1 - \Gamma^{-2})^{1/2}$  and  $t$  is the observer time. If the outflow is beamed and sideways expansion with sound speed takes place, the expression of  $m_{\text{sw}}$  and the half opening angle of the beamed outflow  $\theta$  are respectively given by

$$\frac{dm_{\text{sw}}}{dt} = 2\pi R^2 (1 - \cos\theta) \frac{n m_p \beta c}{1 - \beta}; \quad \frac{d\theta}{dt} = \frac{c_s (\gamma + \sqrt{\gamma^2 - 1})}{R} \quad (3)$$

where  $c_s$  is the sound speed and we use the approximate expression derived by Huang, Dai & Lu (2000), which holds for both the ultra-relativistic and non-relativistic limits.

Assuming that the distribution of the shock-accelerated electrons takes a power-law form with the number density given by  $n(\gamma_e)d\gamma_e = K\gamma_e^{-p}d\gamma_e$  for  $\gamma_m < \gamma_e < \gamma_M$ , the volume emissivity at the frequency  $\nu'$  in the comoving frame of the shocked gas is (Rybicki & Lightman 1979)

$$j_{\nu'} = \frac{\sqrt{3}q^3}{2m_e c^2} \left( \frac{4\pi m_e c \nu'}{3q} \right)^{\frac{1-p}{2}} B_{\perp}^{\frac{p+1}{2}} K F_1(\nu', \nu'_m, \nu'_M), \quad (4)$$

where  $q$  and  $m_e$  are respectively the charge and mass of the electron,  $B_{\perp}$  is the strength of the component of magnetic field perpendicular to the electron velocity,  $\nu'_m$  and  $\nu'_M$  are the characteristic frequencies for electrons with  $\gamma_m$  and  $\gamma_M$  respectively, and

$$F_1(\nu', \nu'_m, \nu'_M) = \int_{\nu'/\nu'_M}^{\nu'/\nu'_m} F(x) x^{(p-3)/2} dx \quad (5)$$

with  $F(x) = x \int_x^{+\infty} K_{5/3}(t) dt$  ( $K_{5/3}(t)$  is the Bessel function). The physical quantities in the pre-shock and post-shock ISM are connected by the jump conditions (Blandford & McKee 1976):  $n' = \frac{\hat{\gamma}+1}{\hat{\gamma}-1} n$ ,  $e' = \frac{\hat{\gamma}+1}{\hat{\gamma}-1} (\Gamma - 1) n m_p c^2$ , where  $e'$  and  $n'$  are the energy and the number densities of the shocked gas in its comoving frame and  $\hat{\gamma}$  is the adiabatic index, a simple interpolation of which between ultra-relativistic and non-relativistic limits,  $\hat{\gamma} = (4\Gamma + 1)/3\Gamma$ , gives a valid approximation for trans-relativistic shocks.

Assuming that shocked electrons and the magnetic field acquire constant fractions ( $\epsilon_e$  and  $\epsilon_B$ ) of the total shock energy, we get  $\gamma_m = \epsilon_e \frac{p-2}{p-1} \frac{m_p}{m_e} (\Gamma - 1)$ ,  $B_{\perp} = \sqrt{8\pi\epsilon_B e'}$  and  $K = (p-1)n'\gamma_m^{p-1}$  for  $p > 2$ . From the spectrum  $F_{\nu} \propto \nu^{-0.74 \pm 0.14}$  of the radio afterglow, we infer that  $p \simeq 2.5$ . It is reasonable to believe that  $\nu'_M$ , in comparison with the radio frequencies, is very large throughout the observations. The observer frequency  $\nu$  relates to the frequency  $\nu'$  in the comoving frame by  $\nu = D\nu'$ , where  $D = 1/(\Gamma(1 - \beta \cos\theta))$  is the Doppler factor. The observed flux density at  $\nu$  is given by

$$F_{\nu} = V_{\text{eff}} D^3 j_{\nu'} / 4\pi d^2 \quad (6)$$

where  $V_{\text{eff}}$  is the effective volume of the post-shock ISM from which the radiation is received by the observer and should be  $V = m_{\text{sw}}/n'm_p\Gamma^2$  for the isotropic case.

Below we shall study the afterglows from the relativistic outflow with  $\Gamma_0 \sim 10$  and the mildly-relativistic one respectively.

### 2.1. Ultra-relativistic outflow

In terms of Eqs(1)-(3), we can obtain the dynamic evolution of the outflow, i.e. we get  $\Gamma(t)$ ,  $R(t)$ ,  $m_{\text{sw}}(t)$  and  $\theta(t)$ . Then using Eq.(4) and the expressions for  $B_{\perp}$ ,  $K$  and  $\gamma_m$ , we can get the evolution of the observed flux with

<sup>1</sup>Numerical calculation with the formula about the synchrotron self-absorption process in Wang, Dai & Lu (2000) shows that the synchrotron self-absorption optical depth at the peak time is several orders of magnitude lower than unity. So, crossing of the self-absorption frequency through the observation band is not a viable explanation for this peak.

<sup>2</sup>We estimate that the synchrotron self-absorption frequency is below  $10^7 \text{ Hz}$  at this time.

time. The radio afterglow of August 1998 flare peaks at about one week after the burst. In the relativistic shock model, it requires that the peak frequency  $\nu_m$  crosses the observation band at the peak time if the synchrotron radiation is optically thin <sup>1</sup>. However, this can hardly be satisfied for a ultra-relativistic outflow with  $\Gamma_0 \sim 10$ , for reasonable values of the shock parameters and the medium density  $n$ . Instead, the model light curves generally peak at  $t < 0.1 \text{ days}$ . This can be clearly seen in Fig. 1, in which we plot the model light curves with different values of  $n$  and  $\epsilon_B$  for both the isotropic and beamed outflow cases and compare them with the observation data. Moreover, the peak flux  $F_{\nu_m}$  are generally much larger than the observed peak flux. The reason can be easily understood from the following analytic estimate for the isotropic outflow case.

At the peak time  $t \sim 10 \text{ days}$  of the radio afterglow, the shock had entered the non-relativistic phase, so the radius of the shock is roughly  $R \simeq (5/2)\beta ct$ . From the condition of energy conservation  $\frac{4}{3}\pi R^3 \frac{\beta^2}{2} n m_p c^2 = E$ , one can get  $\beta = (12E/125\pi c^5 t^3 n m_p)^{1/5} = 0.16 E_{44}^{1/5} t_1^{-3/5} n_0^{-1/5}$ . The magnetic field is  $B = \sqrt{8\pi\epsilon_B e'} = 1.4 \times 10^{-3} \text{ G } \epsilon_{B,-3}^{1/2} E_{44}^{1/5} t_1^{-3/5} n_0^{3/10}$  and the peak frequency and the peak flux are, respectively, given by

$$\nu_m = 3 \times 10^5 \text{ Hz } \left( \frac{\epsilon_e}{0.3} \right)^2 E_{44}^{3/5} t_1^{-9/5} n_0^{-1/10} \epsilon_{B,-3}^{1/2}; \quad (7)$$

$$F_{\nu_m} = \frac{N_e P_{\nu,m}}{4\pi d^2} = 4.4 \times 10^7 \mu\text{Jy } E_{44}^{4/5} t_1^{3/5} n_0^{7/10} \epsilon_{B,-3}^{1/2} \quad (8)$$

where  $N_e$  is the total number of the swept-up electrons and  $P_{\nu,m}$  is the peak spectral power (Sari, Piran & Narayan 1998). It is clearly seen that  $\nu_m$  can hardly be as large as  $\nu_{\text{obs}} = 8.46 \text{ GHz}$  for reasonable shock parameters of  $\epsilon_e$  and  $\epsilon_B$  (e.g. Granot, Piran & Sari 1999; Wijers & Galama 1999; Panaitescu & Kumar 2002) and, furthermore, the peak flux is much larger than detected from the giant flare. Though this analytic estimate is for an isotropic outflow case, the beamed outflow has also this problem as shown in Fig. 1.

Although the ultra-relativistic shock associated with the initial hard spike of the giant flare could not be responsible for the observed radio afterglow, we know from Fig. 1 that its radio afterglow emission should be easily detected at the early time even for the beamed case. The optical afterglow emission from the ultra-relativistic shock is also calculated and shown in Fig. 2. Clearly, early optical afterglow emission can be as bright as  $100 \mu\text{Jy}$  (R band magnitude  $m_R = 19$ ) at  $t \lesssim 0.1 \text{ days}$  for  $\epsilon_B \gtrsim 10^{-3}$ , a reasonable value we know from GRB afterglows. So we urge early follow-up radio-to-optical observations for future SGR giant flares to test the presence of ultra-relativistic outflows. Even at late time  $t \sim 10 \text{ days}$ , the flux at low radio frequencies are intense enough to be detected. At, say,  $\nu = 150 \text{ MHz}$ , the extrapolated flux is  $0.4 \text{ Jy}$  from Eqs(7) and (8) for the typical parameters used.<sup>2</sup> The X-ray afterglow emission from the ultra-relativistic shock is, however, predicted to be lower than the detected bright

and pulsed X-ray afterglow flux (Feroi et al. 2001), which is attributed to the emission from the neutron star surface immediately after the giant flare.

### 2.2. Mildly or sub-relativistic outflow

In §2.1, we showed that the radio afterglow from the 1998 August giant flare of SGR1900+14 is inconsistent with the shock emission from an ultra-relativistic outflow. But a mildly or sub-relativistic forward shock is able to explain the observations, as we show below. The reason is that for a mildly-relativistic or sub-relativistic outflow, it has an enough long period of coasting phase of the outflow, during which the flux increases with time even though  $\nu_m \ll \nu_{\text{obs}}$ . After this coasting phase, the shock decelerates and the flux begins to fade in a power law manner. For an isotropic outflow, the coasting phase lasts about

$$t \simeq \left( \frac{3E_0}{2\pi n m_p \beta_0^5 c^5} \right)^{1/3} = 5 \times 10^5 \text{s} E_{0,44}^{1/3} n_0^{-1/3} \left( \frac{\beta_0}{0.4} \right)^{-5/3} \quad (9)$$

until the mass of the ISM swept up by the blast wave is comparable to the mass of the outflow, where  $\beta_0$  is the initial velocity of the outflow. During the coast phase,  $\beta = \text{const}$  and  $R = \beta t \propto t^1$ . According to Eqs.(4) and (6),  $F_\nu \propto R^3 \beta^{(5p-3)/2} \propto t^3$ . When the mildly or sub-relativistic outflow is decelerated by the swept-up mass, it quickly enters the Sedov phase, during which  $\beta \propto t^{-3/5}$  and  $R \propto t^{2/5}$ . So,  $F_\nu \propto t^{(21-15p)/10} \propto t^{-1.65}$  for  $p = 2.5$ . For a beamed outflow, the coasting phase is similar to the isotropic case, as the expansion is dominated by the cold ejecta during this phase. But during the deceleration phase, the shocked ISM plasma has comparable energy to initial energy  $E_0$  and the sideways expansion may take place, so it is expected that the flux may decay steeper than the isotropic case (e.g. Rhoads 1999; Sari, Piran & Halpern 1999).

The fits with model light curves for mildly or sub-relativistic outflow are presented in Figs. 3 and 4 for different shock parameters. In Fig.3, the isotropic energy is chosen to be  $E_{0,\text{iso}} = 10^{44} \text{erg}$  while in Fig 4 the real energy of the ejecta is chosen to be  $10^{44} \text{erg}$  (its isotropic energy is therefore  $10^{44}/(\theta_j^2/2) \text{erg}$ , where  $\theta_j$  is the initial half opening angle). The beamed outflow model can provide nice fits of the observations for a wide range of shock parameters such as  $n$  and  $E_0$ . In Fig.3, we also present the model light curve for an isotropic outflow (without sideways expansion) denoted with the dashed line. Clearly, it decays too slowly to fit the observations. In all these fits, we used fixed values for  $E$ ,  $p$ ,  $\theta_j$ ,  $\epsilon_e$  and  $n$  with only two free parameters: the initial Lorentz factor  $\Gamma_0$  and  $\epsilon_B$ . We therefore conclude that the mildly or sub-relativistic outflow from the SGR giant flare could provide a plausible explanation for this radio afterglow.

### 3. DISCUSSIONS AND CONCLUSIONS

We have shown that a mildly or sub-relativistic outflow from the SGR could be consistent with this radio afterglow. This outflow is expected to originate from the neutron star crust, accompanying the giant flare. SGRs are now believed to be "magnetars", neutron stars with surface field of order  $10^{14} - 10^{15}$  Gauss or more (Duncan & Thompson 1992; Thompson & Duncan 1995). The en-

ergy that drives the very large ( $\gtrsim 10^{44} \text{erg}$ ) giant flares such as the March 5, 1979 event from SGR 0526-66 and the August 27, 1998 event from SGR 1900+14 is attributed to a sudden large-scale rearrangement of the magnetic field which releases the magnetic energy, while the smaller short repeating bursts ( $E \lesssim 10^{41} \text{erg}$ ) seem to be well explained as being driven by the localized yield of the crust to the magnetic strain. A magnetic field with  $B > (4\pi\phi_{\text{max}}\mu)^{1/2} \sim 2 \times 10^{14} (\phi_{\text{max}}/10^{-3})^{1/2} \text{G}$  can fracture the crust, where  $\mu \sim 10^{31} (\rho/\rho_{\text{nuc}})^{4/3} \text{ergcm}^{-3}$  is the shear modulus of the crust and  $\rho_{\text{nuc}}$  is the nuclear density and  $\phi_{\text{max}}$  is the yield strain of the crust. However, such a patch of crust is too heavy to be able to overcome the binding energy of the neutron star. We expect that only a tiny fraction of the fracturing crust matter can overcome the gravitational binding energy and is able to be accelerated to a mildly-relativistic velocity  $\Gamma_0 - 1 \sim 0.1$  by the released magnetic field energy. Note that the kinetic energy of the mildly-relativistic matter per unit of mass  $(\Gamma_0 - 1)c^2$  is comparable to the binding energy  $GM_{\text{NS}}/R$ , where  $M_{\text{NS}}$  and  $R$  are the mass and radius of the neutron star respectively. Let's denote the amount of matter as  $\Delta m$ , the isotropic kinetic energy  $E_0$  and the real energy of the beamed outflow  $E_r = E_0 \theta_j^2/2$ , where  $\theta_j$  is the beaming angle of this outflow. For  $\Gamma_0 - 1 \sim 0.1$ ,  $\Delta m = 5 \times 10^{22} E_{r,43} \text{g} = 5 \times 10^{23} E_{0,44} \theta_j^2 \text{g}$ . Let the size of this patch of matter be  $\Delta r$ . Because of the insensitivity of  $\Delta r$  on  $\Delta m$  for the outermost crust of neutron star, we estimate  $\Delta r \simeq 0.1 - 0.3 \text{Km}$  for  $E_r \simeq 10^{42} - 10^{44} \text{erg}$ .

Once we know  $\Delta r$ , we can estimate the beaming angle  $\theta_j$  of the outflow when it breaks away from the confinement of the magnetic field. This amount of matter will be vaporized and become plasma near the neutron star surface, which moves out along the open field lines of the magnetar. The initial kinetic energy density of this outflow is  $\epsilon_{k0} = \dot{E}_r/(A_0 \beta_0 c) = 1.6 \times 10^{24} \text{ergcm}^{-3} \dot{E}_{r,43} (\Delta r/2 \text{Km})^{-2} (\beta_0/0.4)^{-1}$ , where  $\dot{E}_r$  is the real kinetic energy luminosity of this outflow,  $\beta_0 c$  is the initial velocity of this outflow and  $A_0$  is the initial sectional area. As the plasma moves out to radial radius  $r$ , the sectional area  $A = \pi(r \sin \theta)^2$ , where  $\theta$  is the angle relative to the magnetic axis. Because the magnetic field lines satisfy  $r \propto \sin^2 \theta$ ,  $A \propto r^3$  for small  $\theta$  and  $\epsilon_k \propto r^{-3}$ . On the other hand, the magnetic field energy density scales with  $r$  as  $\epsilon_B = (B_0^2/8\pi)(r/R)^{-6}$ , where  $B_0$  is the surface magnetic field of the neutron star. When the magnetic field energy density decreases to be comparable to the kinetic energy density, the outflow plasma breaks from the confinement of the magnetic field. This corresponds to a radius  $r_b/R \simeq (B_0^2/8\pi\epsilon_{k0})^{1/3} \simeq 30 B_{0,15}^{2/3} \dot{E}_{r,43}^{-1/3} (\beta_0/0.4)^{1/3} (\Delta r/0.2 \text{Km})^{2/3}$  and  $\theta_b/\theta_0 \simeq (r/R)^{1/2} \simeq 5.5 B_{0,15}^{1/3} \dot{E}_{r,43}^{-1/6} (\beta_0/0.4)^{1/6} (\Delta r/0.2 \text{Km})^{1/3}$ . Note that  $\theta_b$ , which is roughly equal to the beaming angle  $\theta_j$ , is very insensitive to the value of  $\dot{E}_r$ . As the initial opening angle near the neutron star surface  $\theta_0 = \Delta r/R = 0.02(\Delta r/0.2 \text{Km})$ , we estimate the beaming angle of the outflow is  $\theta_j \simeq 0.1 - 0.2$  for typical parameters.

What powers the ejection of this patch of matter? We think that the reconnection of the magnetic field within a region of size  $\Delta r$  during the period of the giant flare will release energy of  $(B^2/8\pi)(\Delta r)^2 V_A \Delta t \sim (B^2/8\pi)(\Delta r)^2 R$ ,

which should be equal to  $\Delta m(\Gamma_0 - 1)c^2$ , where  $B$  is the crust magnetic field,  $V_A$  is the internal Alfvén velocity and  $\Delta t$  is the growth time of the instability causing the giant flare, viz. the duration ( $\Delta t \simeq 0.5\text{s}$ ) of the initial hard spike of the August 27 giant flare (Thompson & Duncan 1995, 2001). The size is therefore estimated to be  $\Delta r \sim 0.2\text{Km}$  and the mass is  $\Delta m \sim 10^{23} B_{15}^2 \text{g}$ , which are in reasonable agreement with the above estimates according to the light curve fits if  $E_0 \sim 10^{45}\text{erg}$  or equally  $E_r \sim 10^{43}\text{erg}$ .

Note that continuing acceleration of electrons at the surface of the neutron star is also a possible mechanism for the radio afterglow, and need further careful investigation in future.

In summary, we studied the afterglow emission from the possible ultra-relativistic outflow and mildly or sub-relativistic outflow accompanying the SGR giant flares. The radio afterglow emission from the August 27 giant

flare of SGR 1900+14 is consistent with a mildly or sub-relativistic outflow, but could not be produced by the forward shock emission from an ultra-relativistic outflow. However, we predict that this ultra-relativistic outflow, suggested to be associated with the hard spike of the giant flare, if exist, should produce bright radio to optical afterglows at the early phase ( $t \lesssim 0.1\text{days}$ ), which can be tested by future observations.

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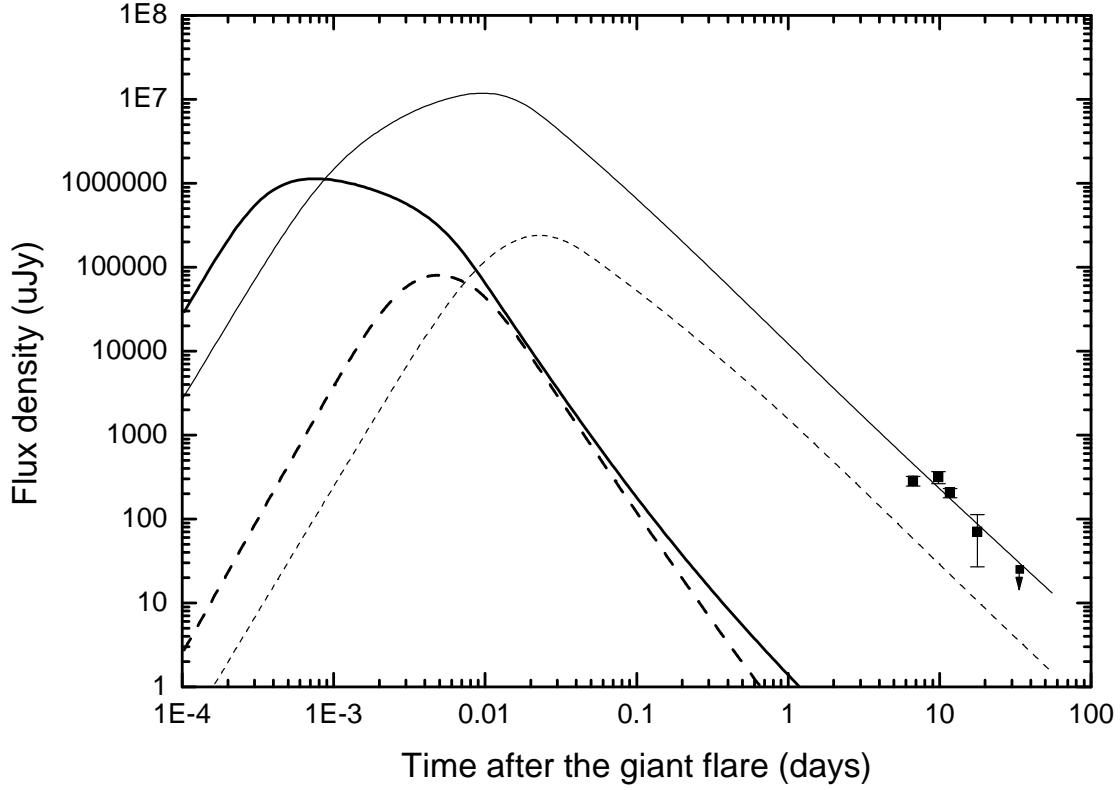


FIG. 1.— Comparison between the model light curves of the afterglows from ultra-relativistic outflows ( $\Gamma_0 = 10$ ) with the observations of the radio flare from the August 27 giant flare of SGR 1900+14 observed at the frequency 8.46GHz. Detections and upper limits for the non-detections, taken from Frail et al. (1999), are indicated by the filled squares and arrows respectively. The *thin* solid and dashed lines represent the afterglows of isotropic outflows expanding into ISM with  $n = 1\text{cm}^{-3}$  and  $n = 0.01\text{cm}^{-3}$  respectively. Other parameters used are  $E_{\text{iso}} = 10^{44}\text{erg}$ ,  $\epsilon_e = 0.3$ ,  $\epsilon_B = 10^{-5}$ . The *thick* solid and dashed lines represent the afterglows of beamed outflows with  $\theta_j = 0.15$  expanding into ISM with  $n = 1\text{cm}^{-3}$  and  $n = 0.01\text{cm}^{-3}$  respectively. Other parameters are the same as thin lines.

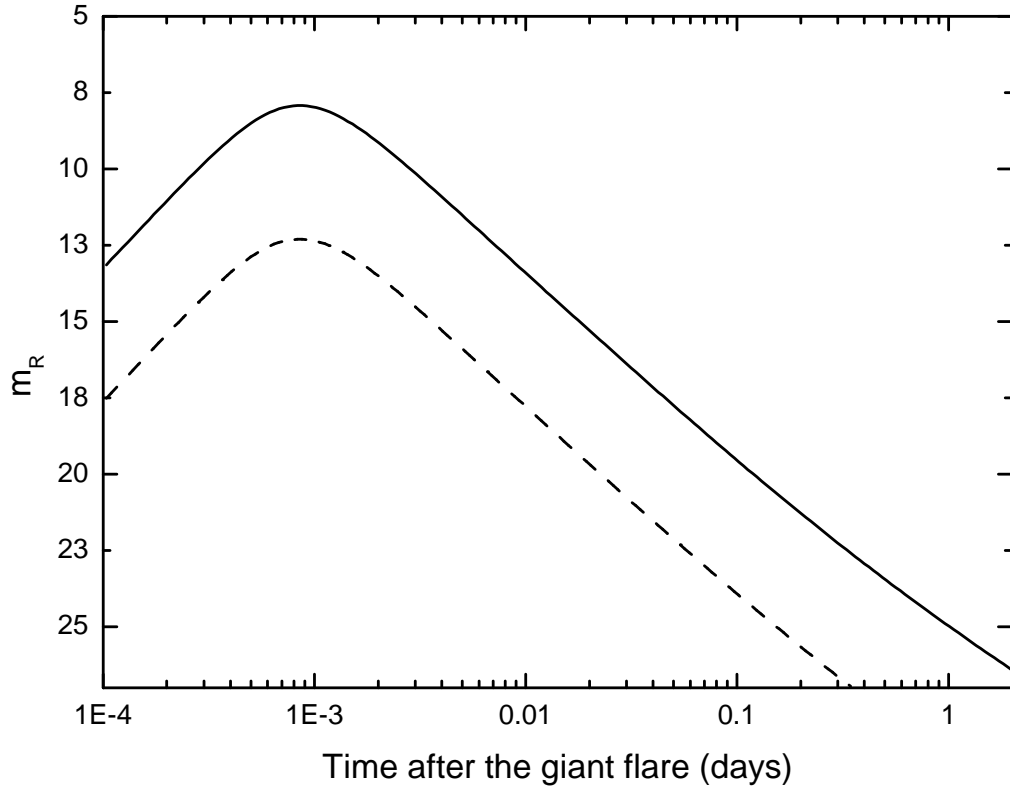


FIG. 2.— Predicted  $R$ -band ( $\nu_R = 4.4 \times 10^{14}$  Hz) optical afterglow light curves of beamed ( $\theta_j = 0.15$ ) and ultra-relativistic ( $\Gamma_0 = 10$ ) outflows from SGR giant flares with  $E_{\text{iso}} = 10^{44}$  erg,  $n = 1 \text{ cm}^{-3}$  and  $\epsilon_e = 0.3$ , but with different values for  $\epsilon_B$ . The solid and dashed lines correspond to  $\epsilon_B = 10^{-3}$  and  $\epsilon_B = 10^{-5}$  respectively.

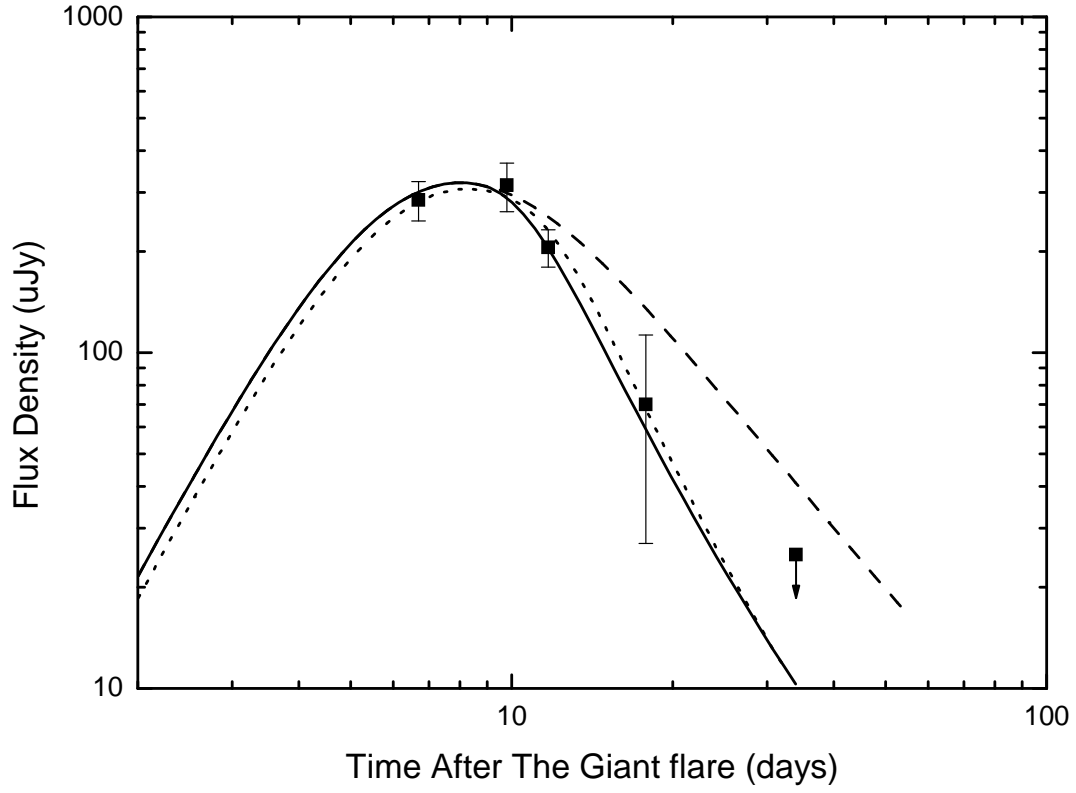


FIG. 3.— Fits of the radio (8.46GHz) flare from August 27 SGR giant flare with afterglow emission from beamed, mildly or sub-relativistic outflow with isotropic energy  $E_{\text{iso}} = 10^{44} \text{erg}$  and  $\theta_j = 0.15$ . The solid and dotted lines correspond to beamed outflows with sideways expansion expanding into ISM with  $n = 1 \text{cm}^{-3}$  and  $n = 0.01 \text{cm}^{-3}$  respectively. Other parameters used are, respectively, ( $\Gamma_0 = 1.033$ ,  $\epsilon_e = 0.3$ ,  $\epsilon_B = 0.008$ ) and ( $\Gamma_0 = 1.13$ ,  $\epsilon_e = 0.3$ ,  $\epsilon_B = 0.03$ ). The dashed line have same parameters as the solid line except that no sideways expansion is considered.

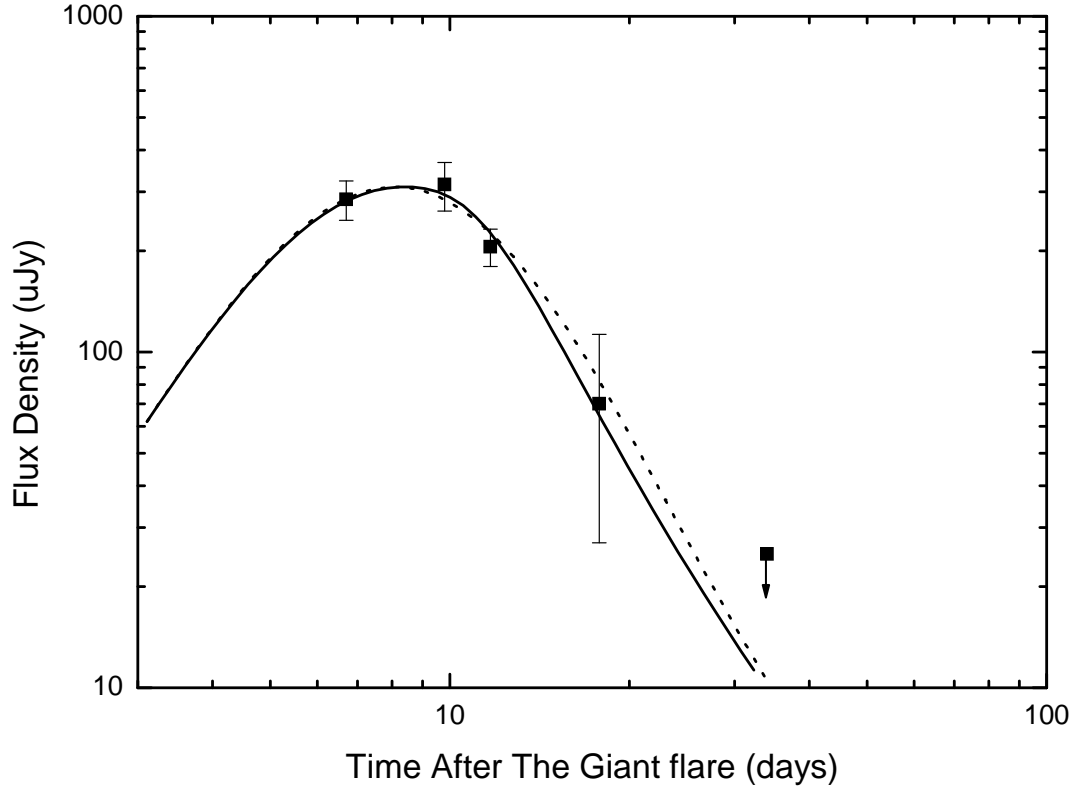


FIG. 4.— Same as Fig. 3 but with  $E_{\text{iso}}\theta_J^2/2 = 10^{44}\text{erg}$ . Parameters used are  $(n = 1\text{cm}^{-3}, \Gamma_0 = 1.13, \epsilon_e = 0.3, \epsilon_B = 3 \times 10^{-6})$  for the solid line and  $(n = 0.01\text{cm}^{-3}, \Gamma_0 = 1.4, \epsilon_e = 0.3, \epsilon_B = 1.5 \times 10^{-5})$  for the dotted line, respectively.